

End-to-end algorithm for hierarchical area searches for long-duration GW sources for GEO 600

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We describe a hierarchical, highly parallel computer algorithm to perform searches for unknown sources of continuous gravitational waves — spinning neutron stars in the Galaxy — over wide areas of the sky and wide frequency bandwidths. We optimize the algorithm for an observing period of 4 months and an available computing power of 20 Gflops, in a search for neutron stars resembling millisecond pulsars. We show that, if we restrict the search to the galactic plane, the method will detect any star whose signal is stronger than 15 times the 1σ noise level of a detector over that search period. Since on grounds of confidence the minimum identifiable signal should be about 10 times noise, our algorithm does only 50% worse than this and runs on a computer with achievable processing speed.

1 Surveys for neutron stars

In this paper we describe progress on developing an efficient computer algorithm to search large areas of the sky for continuous gravitational wave signals from previously unknown sources, most likely spinning neutron stars. The enormous computational cost of processing several months of data by repeatedly applying matched filtering to the data for possible sources lying in each resolved area of the sky is well known.^{1,2} We have previously presented several components of an algorithm that makes use of hierarchical methods to improve search speeds and retain good sensitivity.^{3,4,5} Here we put the components together, estimate the computing requirements of each step (floating-point-operation count), and make a first preliminary optimization of the performance of the search for a given available computer power. We show that the speedup of the method is enough to make area searches practical with the kind of computers that GEO600 may have available to it.

Probably more than 10^8 neutron stars have been formed in the Galaxy in its evolution. Only about 10^3 are known as pulsars or X-ray sources. GEO600 and other gravitational wave projects plan to do directed searches for radiation from those that are known to be spinning

rapidly enough for their radiation to be in the detection window (above 50 Hz for GEO600). But it would also be very interesting to identify new objects by the gravitational waves they emit: if even a small fraction of the unidentified neutron stars in the Galaxy are detectable, we would learn much from them about stellar evolution and the physics of neutron stars. To find such objects requires a blind search.

Recent observations and theoretical work give some reason to believe that some stars may be detectable. The recently discovered *r*-mode instability⁶ will lead to strong radiation from very young stars (~ 1 yr), and a residual radiation may persist for longer times. The youngest neutron star in the Galaxy may be only 40 y old, and if SN1987a contains a neutron star then it will be very young indeed. There are suggestions that low-mass X-ray binaries (LMXB's) are strong gravitational wave emitters⁷, and some of these will probably be targetted by GEO600. But there may be systems where the X-ray emission has turned off relatively recently but the gravitational wave emission continues at some level, and these could turn up in an area search. And of course there may be unexpected reasons for strong emission: the physics of neutron stars is complex and not at all understood.

Another reason for being interested in developing search methods is that they could be applied to other important problems. Detecting radiation from known LMXB's may require searches over large parameter spaces, as almost certainly will the detection of radiation from *r*-mode spindown after a neutron star has been formed in a supernova. Radio searches for pulsars in binary systems have much in common with our problem, and methods developed for gravitational waves could be used in radio astronomy.

These algorithms are of particular interest to the GEO600 project because of the likelihood that it will do a significant amount of observing as a single detector, not in coincidence with other detectors. According to the published plans, GEO600 may be taking data of quality before the large LIGO and VIRGO projects, and it will be unique among the first interferometers in being able to run in narrow-band mode (signal recycling), which will give it better sensitivity to continuous signals than the larger projects would have in the selected bands. Area surveys are likely to be among the first priorities for GEO600's operations when they begin in 2000.

2 Intrinsic difficulty of blind surveys

The computational cost of a blind survey is large because the signal, although predictable, depends on a number of parameters, and there is a very large number of distinguishable sets of values for these parameters. This is primarily because detectors must observe for times of order 10^7 s in order to have reasonable sensitivity. During this time the motion of the Earth produces very substantial phase modulation in the received signal, and the spindown of a neutron star can produce a significant change in its frequency. Since signals can only be found if one can track their phase accurately to within one cycle over the entire observing period, one must search for signals over a huge parameter space.

Consider the effect of phase modulation. The detector is carried by the Earth as it orbits the Sun, and so if it observes a steady source coherently for many months then it effectively synthesizes a gravitational wave telescope with a diameter approaching 2 AU: this is similar to the aperture synthesis common in radio astronomy. The angular resolution of such a telescope is of order $\lambda/2\text{AU}$, or about 0.2 arc seconds. There are about 10^{13} such resolution patches on the sky, each of which impresses a distinguishable pattern of phase modulation onto the signal.

In addition there is spindown, which has been discussed extensively elsewhere.² Young stars in particular can require parametrization of the first three time-derivatives of the period, and this can lead to large (10^{10}) increases in the size of the parameter space.

To search this many parameter sets coherently (i.e. with the optimum sensitivity that can be achieved by matched filtering), each time treating the 2×10^{10} data points that are sampled

in 4 months for observations of up to 1 kHz, is beyond the capacity of any existing or planned computer. In GEO600, the realistic available computing power that can be dedicated exclusively to searches will be of the order of 20 Gflops. Moreover, this computer is likely to be a loosely coupled set of parallel processors (a cluster) rather than a tightly coupled parallel machine. This means that the algorithm must require a minimum of inter-processor communication. Standard signal-processing techniques based on long FFTs do not automatically satisfy this constraint.

3 Hierarchical methods

A solution to this mismatch is to use hierarchical procedures. In general, these involve a step in which candidate sources are selected on the basis of a sub-optimal search, and then they are followed up somehow to test whether they are real or just artifacts of noise. The full data set is never searched at full sensitivity.

The initial selection of candidates inevitably runs the risk that sources will be missed, but in some circumstances this risk is smaller than one might expect. In particular, the large parameter space required for a blind area search implies that there will be many opportunities for noise to masquerade as a signal, so that confidence in detection will require a relatively high signal-to-noise ratio even in optimum filtering. It is conventional in this problem to expect that the best one can hope for is an amplitude SNR of 10.

If this is the case, then one can try to find a method in which the initial sub-optimal is at a level such that a signal of this strength would be likely to get through it. Then very few detectable signals will be lost in such a method.

Our proposed search method consists of the following stages for treating a data set gathered in a total observing time of T_{obs} . The data set is divided into shorter segments of length T_c , for which a full coherent search is performed over the (much smaller) parameter space appropriate to that length of data. The power spectra produced from each such period of time are searched for evidence of a signal whose frequency is changing over the longer time T_{obs} in precisely the pattern expected for some one of the parameter sets, using a method we have adapted from high-energy experimental physics, called the Hough transform.⁴ And finally, candidates are selected at the end of the Hough stage for matched-filtering follow-up over the whole T_{obs} . This uses approximate short-period Fourier transform techniques that we have also developed, to avoid performing large FFTs on parallel computers.³

For a given observing time T_{obs} , which we imagine is set by operational considerations in the experiment, one is free to choose the coherent search length T_c as one wants. The longer it is, the more sensitive will be the search and the more computer power will be required. (Optimum searches would use $T_c = T_{obs}$, but this make impractical demands on computing.) Changing T_c affects the required computing power of the different stages in different ways, and this in turn will depend on the model of the signal: a search will need to define reasonable parameter ranges. We therefore present in this paper the first attempts to *optimize* the search strategy, by fixing the computer power available and choosing T_c in such a way as to give the best search sensitivity.

A different hierarchical approach to the same problem is being developed in the LIGO project.⁸ In this method, the initial coherent transforms are combined using power-spectrum summation, and the final follow-up is done with further power-spectrum summation. It is very important that these two different styles be thoroughly evaluated to decide on their costs and returns for different kinds of problems.

4 Computational cost of the hierarchical procedure

We describe the algorithm in some detail here. It is useful to establish our notation and basic concepts at the beginning. We aim at a total observing time T_{obs} of order 10^7 s, and we expect to do the coherent searches over shorter time T_c of order 1 day. In fact these searches are built from Fourier transforms of even shorter data sets of length $T_s \sim 30$ min, in a way we will describe. At each stage in the calculation one must consider the number of resolvable sets of parameters. We call each set a “patch” in parameter space, and let N_p represent the number of such patches needed in a calculation. Longer data sets need much larger values of N_p . We search for sources in a frequency bandwidth B , with a maximum frequency f_0 .

The hierarchical procedure described above consists of three basic steps:

- I. *matched filtering* on chunks of data of duration T_c much shorter than the total observation time. This stage is often referred to as the *coherent stage*, since it employs the full information of the data, amplitude and phase. It is computationally intensive for the reasons mentioned in the previous section, but it is affordable because the time baseline is short. The outcome of this is a set of N_p FFTs for each data chunk. N_p is the number of patches in parameter space, and every FFT is demodulated according to a patch. This means that if a signal from a patch were present in the data, in the corresponding demodulated time series it would appear as monochromatic and in the corresponding power spectrum the power would be confined to one (two, at the most) frequency bins. Signal to noise ratio in each chunk is $\sqrt{\frac{T_c}{T}}$ times smaller than it would be by matched filtering over T . Each filtered FFT of baseline T_c is actually constructed from a set of shorter baseline (T_s) FFTs, as mentioned in the previous section. The short FFTs should belong to a frequency domain data base which should be constructed as data is acquired. Such data base, as first suggested by Frasca ⁽⁹⁾, should be such that any periodic signal at a given frequency in the data should appear as monochromatic during the observation time baseline. This fixes the maximum length of the baseline at any given frequency. In order to observe a range of frequencies one should ask that the above requirement be fulfilled for the highest frequency of the range. If f_0 is such frequency, then the time baseline for the FFTs of the data base can be chosen to be $T_s = 5.5 \times 10^3 \sqrt{\frac{300 \text{ Hz}}{f_0}}$ s. The use of the short-term frequency database means that area searches can in practice be confined within specific narrow frequency bands: all the data for a particular band can be loaded into the memory of one processor, and negligible data transfers are needed between processors.
- II. the information from the each set of chunks is pieced together by using the *spectra* computed from the demodulated FFTs ^(5, 4, 10). The phase information is lost and this is the reason this stage is said to be *incoherent*. We can show, both analytically and with numerical experiments, that the gain in signal to noise ratio - with respect to that of the previous stage - is the factor $\sqrt[4]{T/T_c}$. At the end of this stage a threshold is set that selects suspect candidates in the parameter space describing the signals one is searching. The strategy that we propose to use for this stage employs a technique which is well known in image processing: the Hough transform (HT, hereafter). It is called “transform” because it is indeed a transformation between the data and the space of the parameters that describe the signal. The outcome of the HT is an histogram in parameter space and significant clustering in a pixel of parameter space indicates “suspect” consistency of data with a signal having the parameters of that pixel. Since the distribution of the number count in each pixel can be known - in fact it is a Poisson distribution - one can associate a false alarm probability, p_{fa} , to each pixel, which measures the significance of

the clustering obtained there. The threshold K is set on this quantity and defines the fraction of parameter space around which a refined search will be performed by step III.

- III a coherent search, as described in step I, is repeated but with the two following differences: 1) the baseline of the FFT is T_{obs} 2) the patches in parameters one corrects for, are the ones “around” the candidates produced by step II. This stage is referred to as the “follow up” stage.

The computational cost of each of these steps can be computed and thus the total computational load may be equated to the available resources. This defines the best acquirable sensitivity and the best choice of algorithm parameters, as we shall show in the following.

Let P indicate the total computation power available, f_0 and B respectively the maximum intrinsic frequency and the band one wants to search over, N_{spin} the total number of sets of values of all the spin-down parameters (which will differ for each step), and A the area of the sky (in steradians) where the search is confined to (e.g. for an all-sky search $A = 4\pi$).

The number of floating point operations χ necessary for each step is the following:

$$\begin{aligned}\chi_I &= 1.2 \times 10^{15} \left(\frac{T_{obs}}{10^7 \text{ s}} \right) \left(\frac{f_0}{300 \text{ Hz}} \right)^2 \left(\frac{B}{300 \text{ Hz}} \right) \left(\frac{T_c}{1 \text{ day}} \right)^2 \left(\frac{A}{4\pi} \right) N_{spin,I}(T_c) \\ \chi_{II} &= 2.5 \times 10^{16} \left(\frac{T_{obs}}{10^7 \text{ s}} \right) \left(\frac{f_0}{300 \text{ Hz}} \right)^2 \left(\frac{B}{300 \text{ Hz}} \right) \left(\frac{T_c}{1 \text{ day}} \right)^2 \left(\frac{A}{4\pi} \right) N_{spin,II}(T_c, T_{obs}) \\ \chi_{III} &= 4.5 \times 10^{23} \left(\frac{T_{obs}}{10^7 \text{ s}} \right)^3 \left(\frac{f_0}{300 \text{ Hz}} \right)^2 \left(\frac{B}{300 \text{ Hz}} \right) \left(\frac{A}{4\pi} \right) N_{spin,III}(T_{obs}) p_{fa}(K)\end{aligned}\quad (1)$$

The cost for step I is the construction and filtering of FFTs over a baseline T_c starting from sets of baseline T_s . The cost for step II is the cost of constructing the HT histograms in the hyperspace of parameters. The cost for step III is the cost of performing a coherent search over the entire observation period around the suspect parameters identified by the previous step. The area around the suspect parameter to be searched is the resolution in parameter space at the end of the HT step (II).

5 Optimization scheme

For a given observation time and fixed computational power P , assuming that data analysis should run at the same speed as data acquisition, the maximum number of operations that one can perform is

$$\chi(T_c, K) = PT_{obs}.$$

We can, though, choose how to distribute the computational load among steps I, II and III. In particular we are free to choose the length of the time baseline T_c and the optimization scheme tells us how to do so in order to achieve the maximum sensitivity. The point is the following: P determines the maximum number of follow-ups that one can perform, i.e. where one should set the threshold K . On the other hand, for a signal of a given amplitude one may determine, given T_c , what p_{fa} it is expected to show up with, after the HT procedure. Conversely, for any given T_c and threshold K , one can say what is the signal to noise ratio $\left(\frac{S}{N}\right)_{min}^2$ at the end of the follow-up stage for the signal which is expected to exceed the threshold 50% of the times. This is what we shall refer to as the smallest expected detectable signal:

$$p_{fa}(S/N) \sim e^{-\frac{1}{2}\sqrt{\frac{T_c}{T_{obs}}}(S/N)^2} \rightarrow \left(\frac{S}{N}\right)_{min}^2 \sim 2\sqrt{\frac{T_{obs}}{T_c}} \ln \frac{1}{K}\quad (2)$$

For given computational power, the optimization scheme consists in determining T_c in such a way that the corresponding expected minimum detectable signal is the smallest possible. Table

5 shows the result of this optimization for the particular case of a search for old (spindown age greater than 10^9 y), fast ($f_0 = 1$ kHz) pulsars using 4 months of data. With 20 Gflops of computing power, it would be possible to detect with 50% efficiency sources with $\text{SNR} \sim 23$, at the end of the follow up. Performance unfortunately is a very slowly varying function of computing power. In fact if the computational resources were increased by a factor 50 the corresponding S/N_{\min} would be reduced only by a factor of 2.

P	T_{obs}	f_0	B	N_{spin}	A	$T_{\text{c/opt}}$	S/N_{\min}
20 Gflops	10^7 s	1000 Hz	500 Hz	1	4π	14 h	~ 23.3
100 Gflops	10^7 s	1000 Hz	500 Hz	1	4π	1.3 days	~ 18.2
1000 Gflops	10^7 s	1000 Hz	500 Hz	1	4π	4 days	~ 12.8
20 Gflops	10^7 s	1000 Hz	500 Hz	1	0.6	2.6 days	~ 14.6
40 Gflops	10^7 s	1000 Hz	500 Hz	1	0.6	3.7 days	~ 13.1
40 Gflops	10^7 s	1000 Hz	25 Hz	1	4π	3.6 days	~ 13.2

Table 1: For different computational resources, and observation time $T_{\text{obs}} = 4$ months, this table shows the performance of the hierarchical algorithm presented in this pape, for searches for old fast pulsars ($f_0 = 1$ kHz and no spindown parameters). The performance is expressed as the signal to noise ratio, S/N_{\min} , necessary for a signal to be detect with a 50% efficiency. Note that this signal to noise ratio refers to a coherent search over the entire T_{obs} .

6 Conclusions

We have described a hierarchical, highly parallel algorithm to perform wide area searches for continuous gravitational wave signals. We have shown that it can be optimized in such a way that, for an observing period of 10^7 s and a 20 Gflops computer, a wide-band wide-area search for old, high-frequency pulsars (typical of millisecond pulsars) can reach an amplitude S/N limit of 23, which is a factor of 2.3 worse than the best one could expect to do with matched filtering (considering the confidence requirements mentioned earlier). If the search is over a more restricted area, such as the galactic plane, then the sensitivity improves to about S/N of 15, only 50% worse than optimum. Restricting a search for millisecond pulsars to the galactic plane is in fact very reasonable: they are formed in binary systems, so they do not acquire the large space velocities of isolated pulsars, and they should accumulate in the galactic plane forever. Radio observations can detect only nearby pulsars, perhaps less than 10% of the galactic population.

These numbers are very encouraging. We have not by any means exhausted the possibilities for speeding up the algorithm, nor have we performed the fullest possible optimization. The final sensitivities can only be better than those quoted here.

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